

Expected masses of merging compact object binaries observed in gravitational waves

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ABSTRACT

We use the well tested StarTrack binary population synthesis code to examine the properties of the population of compact object binaries. We calculate the distribution of masses and mass ratios taking into account weights introduced by observability in gravitational waves during inspiral. We find that in the observability weighted distribution of double neutron star binaries there are two peaks: one for nearly equal mass systems, and one for the systems consisting of a low and a high mass neutron star, $q = 0.6 - 0.7$. The observability weighted distribution of black hole neutron star binaries is concentrated on the systems with the mass ratio $q = 0.3 - 0.5$, while for the double black hole binaries the observability weighted distribution is dominated by the massive, nearly equal mass binaries with $q > 0.7$.

Key words: gravitational waves, stars:binaries

1 INTRODUCTION

We are currently witnessing a large increase in the sensitivity of the gravitational wave observatories. LIGO (Abramovici et al. 1992) is already taking data, the development of VIRGO (Bradaschia et al. 1990) shows great advances, GEO600 (Danzmann & et al. 1992) and TAMA300 (Tsubono 1995) are operational. In the coming years an even more sensitive Advanced LIGO will begin taking data. Out of a number of potential sources of gravitational radiation the most promising are probably mergers of compact object binaries, i.e. binaries consisting of black holes (BH) and/or neutron stars (NS). These are the only sources for which observations in the electromagnetic domain are consistent with emission of gravitational waves. Present efforts to examine data from gravitational wave detectors show that such detections rely heavily on availability of accurate templates. This provides a case for the importance of accurate merger calculations. The data analysis relies on cross correlating the data with a number of templates. Scanning large volume of the parameter space requires using a large number of templates and may hinder detection of a real but low amplitude signal. Any possibility to limit the amount of templates required or to show in which region of the parameter space a detection is most likely may improve chances of seeing the gravitational waves.

Thus it is important to ask the following questions: what are the most likely objects to be observed? what are the most important parameter sets to explore? In order to answer them one needs to investigate the properties of population of compact object binaries. Observations provide us with six double neutron star binaries (Thorsett & Chakrabarty 1999; Burgay et al. 2003). The ra-

dio selected sample of double neutron star binaries is biased toward long lived systems. However, we do not know any black hole neutron star nor double black hole binary. Therefore inferring the properties of the population of compact object binaries solely on the observations of these few systems may lead to erroneous results. A different approach - the binary population synthesis allows to investigate the properties of such systems from a theoretical point of view. Binary population synthesis requires however, a thorough investigation of the systematic uncertainties due to parametrization of various stages of stellar evolution. The population synthesis studies have already been used to estimate the rates and properties of the mergers that can be observed by the gravitational wave observatories (Lipunov et al. 1997b; Fryer et al. 1998; Portegies Zwart & Yungelson 1998; Bethe & Brown 1998; Fryer et al. 1999; Belczyński & Bulik 1999; Bulik et al. 1999; Belczyński et al. 2002c; Nutzman et al. 2004). It has been shown that the observed sample will most likely be dominated by the mergers of double black hole binaries (Lipunov et al. 1997a; Bulik & Belczyński 2003). The distribution of observed chirp masses was found to be a very sensitive indicator of the stellar evolution model while being relatively not sensitive to the star formation rate history and cosmological model (Bulik et al. 2004). A preliminary study of the distribution of mass ratios in compact object binaries was presented by Bulik et al. (2003).

In this paper we use the StarTrack population synthesis code to investigate the distribution of masses and mass ratios in the population of compact object binaries. We use a convention where the mass ratio q in a binary system is defined as the ratio of the lower

Table 1. Population synthesis models. We list the number of coalescing compact object binaries produced in each simulation. For detailed models description see § 2.1 and § 2.2.

Model	Description	N produced
A	standard model described in § 2.1	5761
B1	zero kicks	21535
B7	single Maxwellian with $\sigma = 50 \text{ km s}^{-1}$	17747
B11	single Maxwellian with $\sigma = 500, \text{km s}^{-1}$	2155
B13	Paczynski (1990) kicks with $V_k = 600 \text{ km s}^{-1}$	8270
C	no hyper-critical accretion onto NS/BH in CEs	4798
E1	CE efficiency: $\alpha_{\text{CE}} \times \lambda = 0.1$	894
E2	CE efficiency: $\alpha_{\text{CE}} \times \lambda = 0.5$	3489
E3	CE efficiency: $\alpha_{\text{CE}} \times \lambda = 2$	8504
F1	mass fraction accreted in non-cons. MT: $f_a = 0.1$	2483
F2	mass fraction accreted in non-cons. MT: $f_a = 1$	4644
G1	wind decreased by $f_{\text{wind}} = 0.5$	9395
G2	wind changed by $f_{\text{wind}} = 2$	5517
J	primary mass: $\propto M_1^{-2.35}$	8220
L1	angular momentum of material lost in non-cons. MT: $j = 0.5$	6660
L2	angular momentum of material lost in non-cons. MT: $j = 2.0$	5547
M1	initial mass ratio distribution: $\Phi(q) \propto q^{-2.7}$	852
M2	initial mass ratio distribution: $\Phi(q) \propto q^3$	11225
O	partial fall back for $5.0 < M_{\text{CO}} < 14.0 M_{\odot}$	4116
S	all systems formed in circular orbits	4667
Z1	metallicity: $Z = 0.01$	5199
Z2	metallicity: $Z = 0.0001$	7074

mass component to the higher mass one and therefore is always less than unity. In section 2 we shortly describe the code, and demonstrate the difference between the volume limited and flux limited distributions of masses of compact object binaries. We present the results in section 3, and conclusions in section 4.

2 CALCULATIONS

We are using the *StarTrack* binary evolution code described in detail by Belczynski et al. (2002c). The code is well tested and has been used in various astrophysical applications: analysis of gamma ray-burst progenitors (Belczynski et al. 2002b), tracing of evolutionary history of individual binaries (Belczynski & Bulik 2002), investigation of the mass spectra of compact objects (Belczynski et al. 2002a). *StarTrack* population synthesis code was specifically designed to calculate the merger rates and physical properties of compact object binaries. It was compared with a several other codes (e.g., Lipunov et al. (1997b), Portegies Zwart & Yungelson (1998) de Donder & Vanbeveren (1998) , Fryer et al. (1999), Nelemans et al. (2001)) . The comparisons showed some differences, however they were understood within the different model assumptions. Since *StarTrack* was designed to deal mostly with systems containing NSs and BHs, our input physics was updated and revised as compared to the other codes in respect to compact object formation. As a result we have recognized new evolutionary NS-NS formation scenarios, and we have shown that massive stellar BH may dominate the population of double compact objects ($\sim 10 M_{\odot}$) observed in gravitational waves. Last, but not least, our predictions of NS-NS Galactic coalescence rates (Belczynski et al. 2002c) are in good agreement with the most recent constraints obtained from the observed sample of these systems (Kalogera et al. 2001; Kim et al. 2003).

2.1 Standard Model

Within the code the evolution of single stars is parametrized by the modified formulae of Hurley et al. (2000). The single star evolution includes such stages as the main sequence, evolution on the Hertzsprung gap, red giant branch, core helium burning, asymptotic giant branch, and evolution of helium stars. Two major modifications include low mass helium star evolution and calculation of compact object masses. In particular, following a number of studies (Delgado & Thomas 1981; Habets 1987; Avila-Reese 1993; Woosley et al. 1995) we allow the low mass helium stars ($\leq 4 M_{\odot}$) to develop deep convective envelopes. Presence of convective envelope plays an important role in the behavior of the donor star in the Roche lobe overflow event, and may eventually lead to the development of dynamical instability and common envelope (CE) evolution, and possible tightening of the binary orbit.

The original Hurley et al. (2000) formulae are used to calculate the final CO core mass of a given compact object progenitor at the time of supernova/core collapse event. We use Woosley (1986) stellar models to obtain mass of the final FeNi core corresponding to a given CO core mass. The FeNi core is collapsed to form proto-neutron star, and then we use the results of core collapse hydrodynamical calculation of Fryer (1999) to calculate the amount of fall back material and the final mass of the newly formed compact object. We use the following algorithm to derive the masses of a newly formed compact object M_{rem} :

$$M_{\text{rem}} = \begin{cases} M_{\text{FeNi}} & M_{\text{CO}} \leq 5 M_{\odot} \\ M_{\text{FeNi}} + f_{\text{fb}}(M - M_{\text{FeNi}}) & 5 < M_{\text{CO}} < 7.6 \\ M & M_{\text{CO}} \geq 7.6 M_{\odot} \end{cases} \quad (1)$$

where M_{FeNi} is the mass of the FeNi core, M_{CO} is the mass of the CO core, M is the total mass of the star prior to the explosion, and f_{fb} is the fall back factor, $0 < f_{\text{fb}} < 1$ depending on the mass of the star. This simple formula represents well the results of detailed numerical calculations. We do verify the sensitivity of our

results to changes in the particular numerical values in equation 1, see e.g. model O below. Varying stellar evolution parameters like the strength of winds, or metallicity leads to different core masses for a star of given initial mass and also alters the initial final mass relation for single stars. We find that NSs are formed without a significant amount of fall back material, while BHs are formed either directly (prompt collapse of a massive star) or through partial fall-back of material onto proto-neutron star.

The binary evolution takes into account orbit changes due to wind mass loss, and tidal interactions. Wind mass loss rates are adopted from Hurley et al. (2000) and they depend on the stellar parameters of mass losing component (its composition, mass and evolutionary stage). Specific mass loss rates are adopted for naked helium stars, the luminous blue variables and pulsating stars. In stable mass transfer (MT) calculations we allow for non-conservative evolution. We assume that part (f_a) of the transferred material to accreted onto companion star, while rest is ejected from the system with the specific angular momentum (j , expressed in the units of the binary angular momentum). In our standard model we adopt: $f_a = 0.5$ and $j = 1$. If the RLOF episode is dynamically unstable, we follow the spiral in through the common envelope phase. If system avoids the merger, we calculate the final orbital separation using standard energy conservation based prescription of (Webbink 1984). The evolution through CE phase depends crucially on the efficiency of the orbital energy input into the donor envelope (α_{ce}) and the specific binding energy of the envelope (λ). Only the product of these two largely uncertain quantities enter the calculation, and we use $\alpha_{ce} \times \lambda = 1$ in the standard model, however we also check the sensitivity of the results to this parameter. During the CE spiral in we allow for hyper critical accretion onto NSs and BHs (e.g., Blondin (1986), Chevalier (1989), Chevalier (1993), Brown (1995)). As a result several tenths of solar mass may be accreted onto the compact object, and in particular the top heavy NSs may collapse and form BHs. The full description of the hyper critical accretion treatment is given in the appendix of Belczynski et al. (2002c).

Supernovae explosions are treated in detail. The explosion takes place at the randomly selected place on the orbit. We allow for explosions on the eccentric orbits, for uncircularized systems. We take into account the instantaneous mass and angular momentum loss from the binary system. Also a natal kick is added to the orbital velocity of the newly born compact object to account for the SN asymmetry. Kicks are selected from the bimodal distributions of Cordes & Chernoff (1998), a weighted sum of two Maxwellians, one with $\sigma = 175 \text{ km s}^{-1}$ (80%) and the second with $\sigma = 700 \text{ km s}^{-1}$ (20%). A binary is either disrupted in the explosions, in which case we stop the evolution, or if it survives we follow the evolution on the new binary orbit.

The initial mass of the primary M_{zams}^1 is drawn from a power law initial mass function (IMF) distribution $\propto M^{-2.7}$ (Scalo 1986) within the range $8\text{-}100 M_\odot$. The secondary mass is obtained as $M_{zams}^2 = qM_{zams}^1$, where q is the mass ratio and is drawn from a flat distribution (Kuiper 1935). We allow for eccentric initial orbits, and the eccentricities are drawn from a thermal distribution $\propto e$ (Heggie 1975; Duquennoy & Mayor 1991). Finally, the orbital separation distribution is taken to be flat in $\log a$ (Abt 1983), and separation are chosen from few (so the stars are not formed at the contact configuration) up to maximum of 10^5 solar radii. We evolve our stars for a maximum $T_{Hubble} = 15$ Gyrs. The evolutionary model described above is chosen as our reference (standard) model and marked with latter "A" on the following figures and tables.

2.2 Parameter Study

In order to asses robustness of the results we investigate 20 extra different models of stellar evolution, where we vary the parameters describing various stages of stellar and binary evolution. The models used are listed in Table 1. The range of models represents the current state of knowledge and uncertainties about the binary evolution. All models are calculated with 2×10^6 initial binaries each.

In models marked with letter "B" we vary the distribution of natal kicks compact objects receive when they are formed. This is rather uncertain part of evolutionary model as we still do not know the mechanism behind the SN/core collapse asymmetry (Buras et al. 2003). Therefore, we change the kicks quite drastically, from rather non-realistic model with no kicks (B1) to the very high kicks of model B11. The higher the kicks, the less compact object binaries we form, since the higher kicks tend to disrupt the progenitor systems. This is one of the most important parameter as far as the number of compact object binaries are concerned (close to an order of magnitude change).

Since the CE evolution is another highly uncertain part of our evolutionary scheme, in models "E" we change the efficiency with which orbital energy is transformed into the unbinding the envelope of the donor star, while in model "C" we turn off the accretion onto compact objects during that very short lived phase. In models with small CE efficiency (E1-2) it is found that the number of compact object binaries are significantly reduced. This is due to the fact, that many binaries, evolving through the CE phase, will merge, thus aborting compact object binary formation. On the other hand, increase of the efficiency (E3) or shutting down the accretion at CE phase do not play a very important role on the number of formed compact object binaries.

In models "F" and "L" we consider the results for different treatment of the stable MT phases. In particular, in model F2 we consider the case of conservative evolution (all mass and angular momentum transferred to the companion). Change of the MT mode from the non-conservative (standard model) to conservative evolution (F2), does not change the numbers by much. The model F1 with highly non-conservative evolution, decrease the numbers of formed compact object binaries rather significantly, but it is rather improbable, since the estimated material loss is probably not as high as assumed in model F1 (Meurs & van den Heuvel 1989). Since we have adopted quite a large value for specific angular momentum of lost material in non-conservative MT episodes, we should really concern only the model L1 with lowered specific momentum leaving the binary. As we see from Table 1, the numbers in model L1 are almost the same as for standard model. We thus conclude, that the treatment of non-conservative MT phases does not have a great influence on the number of compact object binaries.

Winds of massive stars may play an important role on the population of compact objects. In model G1 we decrease all the wind mass loss rates by factor of 2. The smaller the winds, the more massive compact objects formed, and more BHs are formed as compared to NSs. However, the total number of compact object binaries is basically unchanged in this model.

The flatter IMF slope of model J increases slightly (as slope was not changed by much) the number of heavy stars (progenitors of compact objects) and thus leads to a slight increase in the number of compact object binaries. The IMF slope for massive stars is rather well determined (Kroupa & Weidner 2003), and as expected the small change does not affect the population.

Change of initial mass ratio distribution may have severe ef-

fect on the numbers of compact object binaries (models M1-2). In model M1 most of the progenitor systems are formed with extremely small mass ratios. Therefore, once they reach first MT phase, it is usually dynamically unstable, leading to spiral in and merger of components, aborting the formation of compact object binary. This explains the small number of formed binaries in model M1, and warrants survival of the systems (large numbers of compact object binaries) with rather equal masses in model M2. Since the initial mass ratio distribution is not easily measurable and constrained, models M1-2, although rather extreme, should be taken account in further analysis.

In the end, several other models, do not have much influence on the production efficiency of compact object binaries. These include models with different metallicities (Z1-2), different assumption on initial eccentricities (S) and finally the model in which we change the regime of the fall back in formation of compact object binaries (O).

The masses of the compact objects are strongly affected in some of the models. In particular the masses change quite drastically in models G1 and G2 where the stellar winds are changed. Decreasing the stellar winds (model G1) allows the star to develop more massive cores and consequently leads to higher masses of the compact objects formed. A similar effect is connected with decreasing the metallicity (models Z1 and Z2), since lower metallicity stars have smaller winds. On the other hand within the model G2, where the winds are artificially increased, massive stars do not have the time to develop massive cores and no compact objects above $3 M_{\odot}$ are produced. The masses of the compact objects are also affected by varying parameters in equation 1. In model O we increase the upper bound of the fall-back range to $14 M_{\odot}$. this leads to smaller masses of the black holes produced as even for high mass cores some fraction of the mass is still expelled. The population of compact object binaries in model E1, with reduced common envelope efficiency contains additional systems with massive black holes. Within model C we turn off the hypercritical accretion onto compact objects in common envelope events. This primarily influences the masses of neutron stars and low mass black holes as they have no possibility to increase significantly.

2.3 Distributions of masses and mass ratios

In the output we note the masses of the compact objects in each binary and the lifetimes: the stellar lifetime from the formation at the zero age main sequence to formation of a double compact object, and the lifetime as a double compact object binary until it merges due to gravitational wave emission. We denote the sum of the two lifetimes as the total lifetime of the binary T .

In this calculation we assume for simplicity that the space is Euclidean. We denote the masses of the components in each binary m_1^i and m_2^i , the mass ratio is $q = m_1^i/m_2^i < 1$. The formation rate of compact object binaries with a given mass ratio q , the mass of the primary m_2 (the greater of the two masses), and the lifetime T at a given cosmic time t is

$$\frac{dF(m_2, q, T, t)}{dm_2 dq dT} = \frac{S(t) f_{sim}}{\langle M_* \rangle N_{tot}} \times \sum_{i=1}^{N_{CCOB}} \delta(m_2 - m_2^i) \delta(q - q^i) \delta(T - T^i), \quad (2)$$

where $S(t)$ is the star formation rate at the time t , f_{sim} is the fraction of stars out of a total population that we simulate, $\langle M_* \rangle$ is the average mass of a binary in the stellar population, and N_{CCOB}

is the number of coalescing compact object binaries formed in a simulation of N_{tot} binaries. Our aim is to calculate the observed merger rate by an observer on Earth at present, which we denote as t_0 . The coalescence rate at a distance r from the observer is then given by

$$\frac{df_{coal}(r)}{dm_2 dq} = \int dt' \frac{dF(m_2, q, t', t_0 - r/c - t')}{dm_2 dq dT}. \quad (3)$$

Inserting equation 3 into 2 we obtain

$$\begin{aligned} \frac{df_{coal}(r)}{dm_2 dq} &= \frac{f_{sim}}{\langle M_* \rangle} N_{tot}^{-1} \times \\ &\times \sum_{i=1}^{N_{CCOB}} \delta(m_2 - m_2^i) \delta(q - q^i) S(t_0 - r/c - T^i). \end{aligned} \quad (4)$$

The observed rate is obtained by integrating equation 4 over the volume in which the binaries are observable

$$\frac{dR}{dm_2 dq} = \int_{V(m_2, q)} dV \frac{df_{coal}(r)}{dm_2 dq} \quad (5)$$

We note that for a constant star formation rate the lifetimes of the binaries do not enter the observed rate.

We first calculate a volume limited distribution of masses, i.e. we assume that all binaries coalescing in a given volume V are observable regardless of the mass m_2 and the mass ratio q . This corresponds to observing for example the entire population of a given galaxy or a galaxy cluster. Here we also assume that the star formation history was constant. The normalized volume limited distribution of masses and mass ratios is

$$P(q, m_2) = N_{CCOB}^{-1} \sum_{i=1}^{N_{CCOB}} \delta(q - q^i) \delta(m_2 - m_2^i), \quad (6)$$

In the case of realistic detectors the volume of integration in equation 5 will depend on m_2 and q . Our calculation of the distribution of masses and mass ratio relevant for detecting merging binaries with gravitational waves follows the calculations presented earlier in (Bulik & Belczyński 2003). Here again we assume that the star formation rate was flat and that the Universe is Euclidean and uniformly filled with stars. The signal to noise in high frequency gravitational wave detectors from an inspiral of a stellar mass binary is given by (Finn & Chernoff 1993; Bonazzola & March 1994; Flanagan & Hughes 1998)

$$(S/N) = \frac{A_i}{d} \left[\frac{\mathcal{M}}{M_{\odot}} \right]^{5/6}, \quad (7)$$

where $\mathcal{M} = (m_1 m_2)^{0.6} (m_1 + m_2)^{-0.2}$ is the chirp mass, d is the distance, and the A_i depends on the details of a particular detector. Thus a coalescence of a binary with a chirp mass \mathcal{M}^i will be visible up to a distance proportional to $(\mathcal{M}^i)^{5/6}$ and the volume of integration in equation 5 will be $V^i \propto (\mathcal{M}^i)^{5/2}$. The observability weighted distribution of masses and mass ratios is therefore

$$P_{obs}(q, m_2) = K^{-1} \sum V_i \delta(q - q^i) \delta(m_2 - m_2^i), \quad (8)$$

where $K = \sum_i V_i$. Again the lifetimes of the binaries do not enter the weights in equation 8 because of assumption of constant star formation rate history. This distribution is more realistic as it corresponds to a case of an instrument with the sensitivity allowing it to detect binary coalescences in a large ensemble of galaxies. Relaxation of the assumption of constant star formation and taking into account a realistic cosmological model has been discussed in

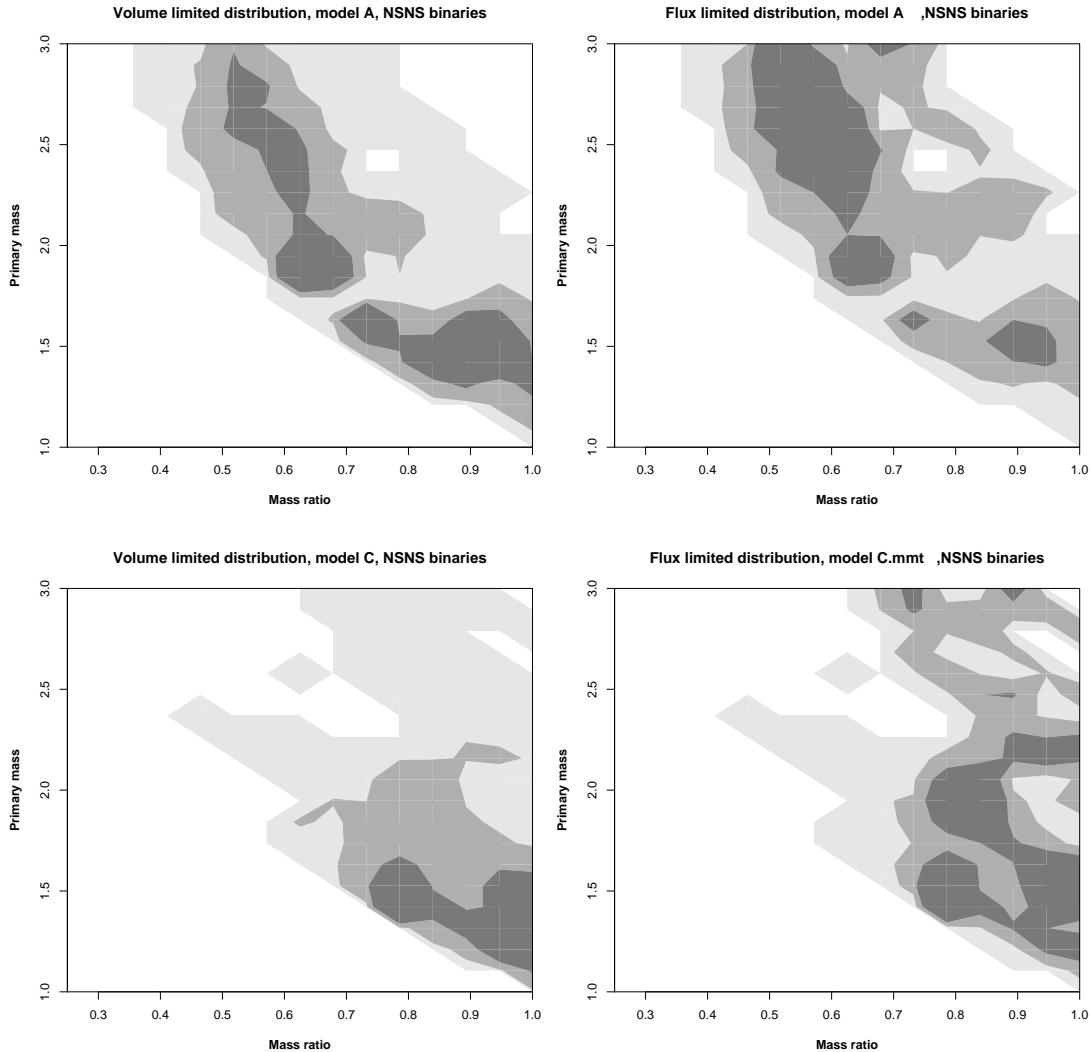


Figure 1. The volume limited and observability weighted distributions of parameters of compact NSNS binaries obtained within model A - top panel and model C - bottom panel. The region in dark gray encompasses 68% of the systems, the medium gray corresponds to 95%, and the light gray corresponds to all binaries in the simulation.

(Bulik et al. 2004) for the case of the distribution of observed chirp masses and was shown not to be significant.

3 RESULTS

In the following we will assume that the maximum mass of a neutron star is $3 M_{\odot}$. All objects above this value will be considered as black holes. In our simulations the minimum mass of a neutron star is $1.2 M_{\odot}$. Thus we can classify all binaries as double neutron star (NSNS), black hole neutron star (BHNS), or double black hole binaries (BHBH). These three categories will be analyzed separately. We will present the distributions of binary parameters in the space spanned by the mass of the primary (the more massive component of a compact object binary) and the mass ratio q .

3.1 NSNS binaries

We present the volume limited distributions of q and m_2 in obtained in the framework of model A in the top panels of Figure 1. The volume limited distribution for the case of NSNS systems (top left panel) exhibits a peak for the systems with nearly equal masses just above the minimal mass of a neutron star. This roughly corresponds with the observations of pulsars where most systems have similar masses. There is however, a long tail in the distribution extending to systems with large mass of the primary, and low mass ratio. these are systems consisting of a neutron star with a mass near the maximum value and a companion neutron star with a low mass. Such systems have chirp mass about 1.5 times larger than the binaries from the above mentioned peak. Therefore in the flux limited distribution - top right panel in Figure 1 -these low mass ratio systems are showing up more prominently. It is however possible that the maximum mass of neutron star is lower than $3 M_{\odot}$ and some of the systems shown here harbor low mass black holes rather than neutron stars. However even for the maximum mass of a neutron star

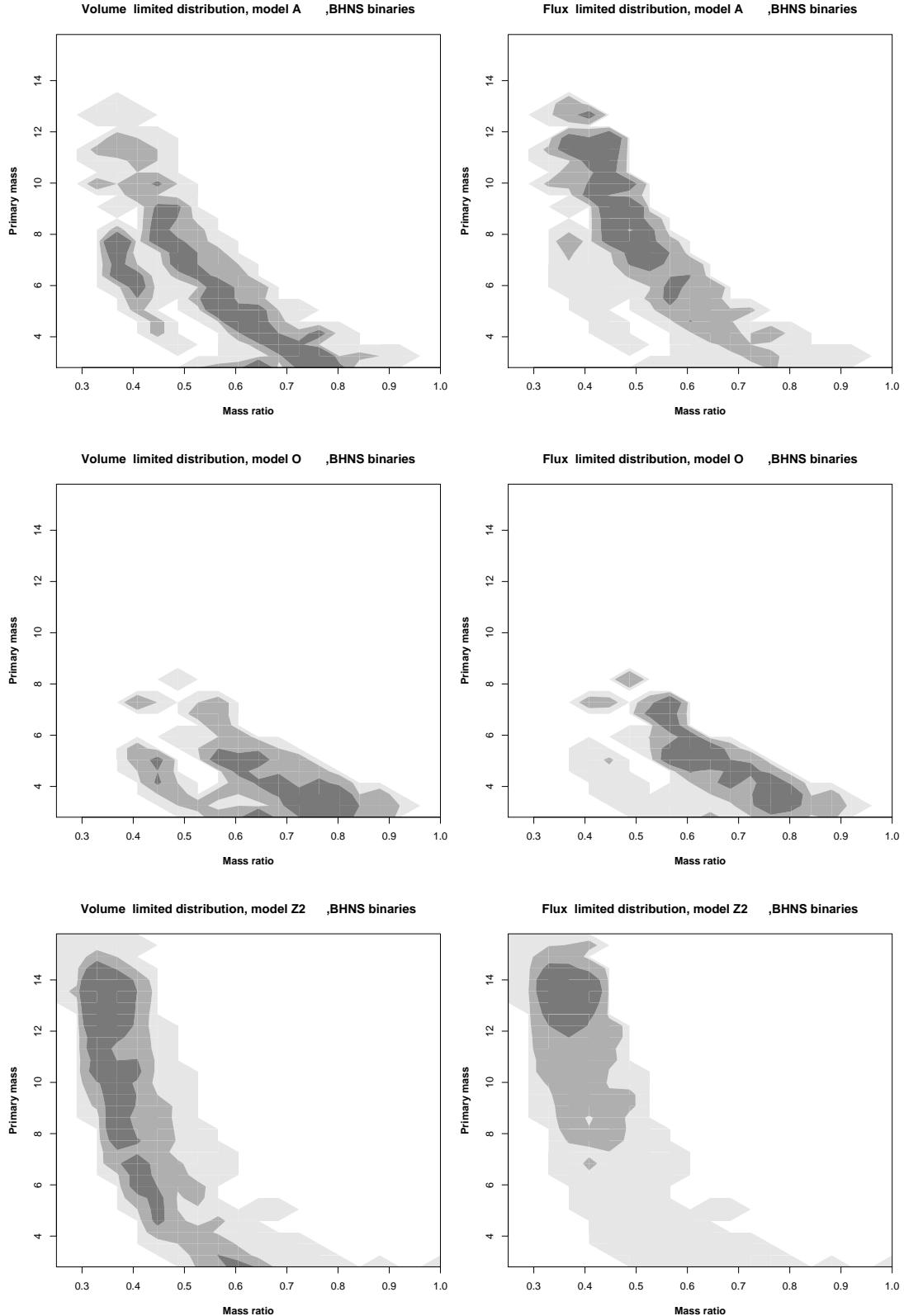


Figure 2. The volume limited and observability weighted distributions of the parameters of BHNS binaries within model A (top panel), model O (middle panel) and model Z2 (bottom panel). The region in dark gray encompasses 68% of the systems, the medium gray corresponds to 95%, and the light gray corresponds to all binaries in the simulation.

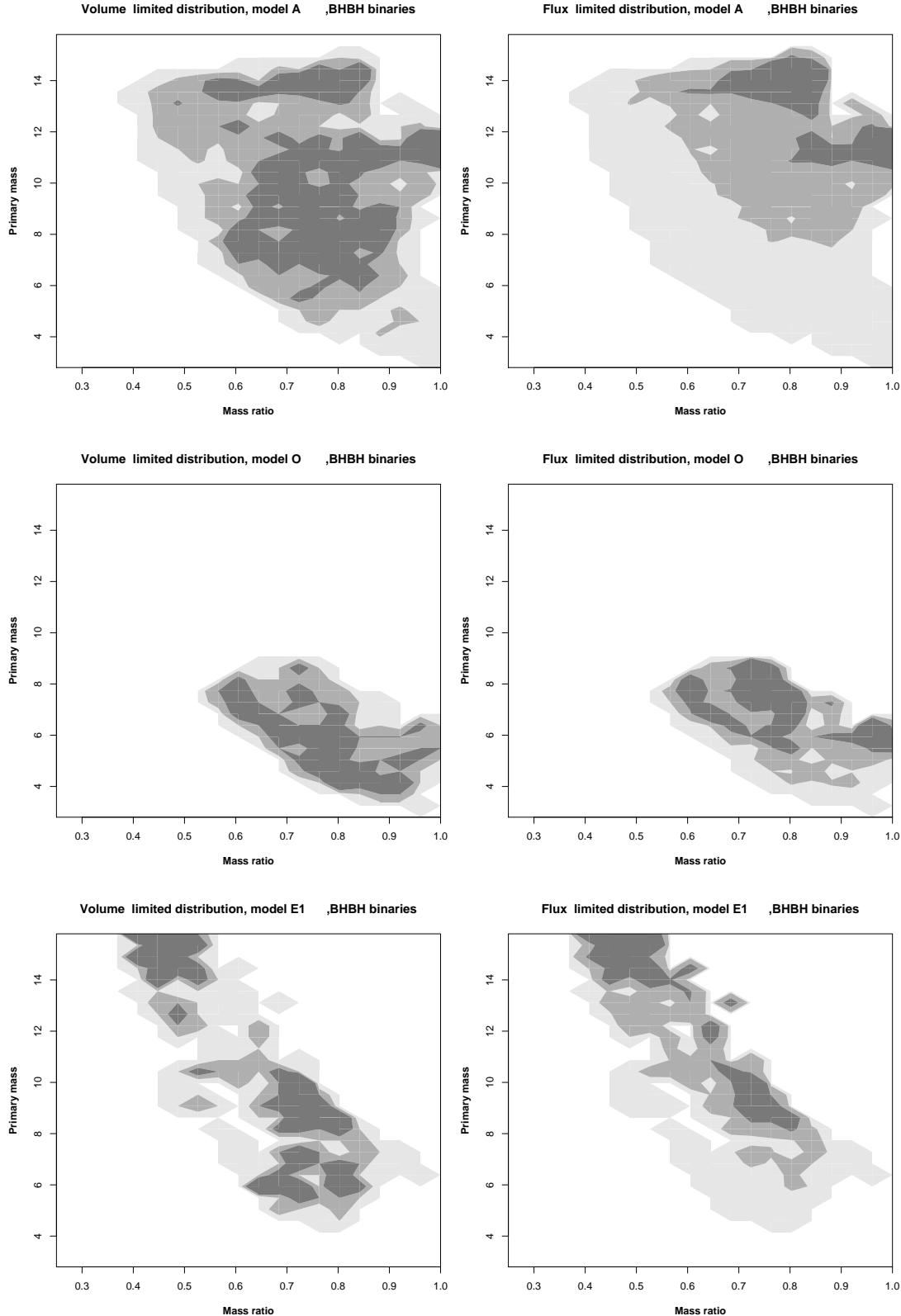


Figure 3. The volume limited and observability weighted distributions of the parameters of BHBH binaries within model A (top panel), and model O (bottom panel). The region in dark gray encompasses 68% of the systems, the medium gray corresponds to 95%, and the light gray corresponds to all binaries in the simulation.

of $2.0 M_{\odot}$ there is still a large fraction of low q systems in the flux limited distribution. We have examined all 21 model listed in Table 1 and nearly all the models show a similar pattern in both the volume limited and flux limited distributions. There is one exception - model C - for which we present the relevant distributions in the bottom panels of Figure 1. In this model we turn off the possibility of hypercritical accretion onto a compact object in common envelope phase. This effectively shuts off the possibility of increasing significantly the mass of a neutron star through accretion. Therefore in the volume limited distribution there is quite a large concentration of systems with both masses below $1.5 - 1.6 M_{\odot}$, and little number of binaries with low mass ratios. Consequently in the flux limited distribution the binaries with low mass ratio are nearly absent in contrast to model A. However, the 68% contour includes systems with $q > 0.75$ and $m_2 \approx 2.0 M_{\odot}$, as well as some binaries with $q \approx 0.75$ and the mass of the primary close to the maximal mass of a neutron star in our model.

3.2 BHNS binaries

We present the volume limited and the flux limited distributions of q and m_2 in Figure 2. The top panel of Figure 2 corresponds to the standard model A. The volume limited distribution shows a large number of binaries along a stripe stretching from $q \approx 0.7$ and $m_2 \approx 4 M_{\odot}$ to $q \approx 0.2$ and $m_2 \approx 10 M_{\odot}$. Systems above and to the right of this stripe would be classified as BHBH binaries. The volume limited distribution is dominated by binaries with low mass black holes. In the flux limited distribution the binaries with higher mass black holes and low mass ratios start to play an important role. This is due to the balance between the falling mass function and the increase in the chirp mass with increase of the mass of the primary. For nearly all models this leads to dominance of binaries where the black hole primary has a mass between 6 and $12 M_{\odot}$, and the mass ratio is somewhere from 0.3 for the most massive black holes to 0.5 for the moderate mass ones.

After examining the 21 models of Table 1 in the case of BHNS binaries one can distinguish two other classes of models with different distributions of binary parameters. The first class consists of models , L2, M2, and O. For this models we show a representative case (model O) in the middle panel of Figure 2. The common characteristic of this class of models is that the population of black holes in binaries lacks the very massive ones for various reasons. In the case of the model O shown in Figure 2 this is because we increase the range of masses of fall-back formation of black holes. Within the models E1, L2, M2 such binaries have a smaller chance of formation because of altering the treatment of the mass transfer events. in consequence the is very little low mass ration BHNS binaries and the volume limited distribution is very similar to the flux limited one.

A separate class of consists of models E1, Z2 and G1. In their case high mass black holes are easily formed, because of decrease of strength of stellar winds (Z2 and G1). Model E1 favors survival of systems with high mass first born stars and leads to effective production of extreme mass ratio compact object binaries. We present the two distributions for the case of model Z2 in the bottom panel of Figure 2. Here the volume limited distribution is dominated by a nearly vertical stripe at $q \approx 0.2$. In the flux limited distribution the binaries with high mass primaries $m_2 \approx 12 M_{\odot}$ and $0.1 < q < 0.3$ are dominant simply because of their high chirp mass.

3.3 BHBH binaries

The case of BHBH binaries is presented in Figure 3. The top panel corresponds to the model A. The volume limited distribution fills more or less uniformly the region allowed for the BHBH binaries. In the flux limited distribution there is a preference for the high mass ratio (nearly equal mass) and high mass ratio systems, i.e. these filling the top right corner of the plot. Thus the flux limited distribution is dominated by the systems with $q < 0.6$ and m_2 near the maximum mass produced in a given model. All models seem to follow this general trend.

In models O the maximum mass of a black hole is decreased. We present the results of a calculation using model O, in the middle panel of Figure 3. In these case the volume limited distribution also fills nearly uniformly the region allowed for black holes. However because of it smaller size the range of chirp masses for given mass ratio is not as large and the flux limited distribution is only slightly shifted to higher masses in respect to the volume limited one.

Another special case - model E1 - is presented in the bottom panel of Figure 3. Here because of lowered CE efficiency formation of equal mass compact object binaries is favored, while extreme mass binaries are preferentially formed. The volume limited distribution in this case is dominated by systems with mass ratio in the range $0.4 < q < 0.6$ and a tail extending to $q = 0.8$.

Models of the class that favors production of massive black holes (Z2, G1) do not lead to qualitatively different results than model A. In these models the flux limited distributions tend to concentrate around binaries with higher mass ratio, $q > 0.7$ and higher total masses than in the standard model A.

4 CONCLUSIONS

We have calculated the expected distributions of masses and mass ratios of compact object binaries to be observed in gravitational waves. The results are based on the Star Track binary population synthesis code. For most of the models the observability weighted distribution of double neutron star systems has two peaks: one with the mass ratio almost unity and both masses near the smallest mass allowed for neutron stars, and another with small mass ratio, consisting of stars with the mass near the maximum mass of a neutron star in a binary with a star close to the minimum mass. The reality of this second peak depends on the assumed maximum mass of a neutron star: the lower the maximum mass of a neutron star the smaller the small mass ratio peak. The distribution of black hole neutron star binaries peaks at mass ratios between 0.3 and 0.5. The bulk of observed double black hole binaries has mass ratios above 0.7. We have shown that these results are rather generic and depend weakly on the choice of a particular model of stellar evolution.

The crucial parameter determining the shape of the distribution of the observed NSNS binaries is inclusion of the hypercritical accretion onto compact objects in common envelope events. In the case of BHNS and BHBH binaries the most important parameters are these that alter the masses of the black holes in such binaries, and the common envelope efficiency. The masses may be altered either due to the mechanism of compact object formation in supernova explosions, or due to particular treatment of mass transfer events. The distribution of the BHNS binary parameters is most sensitive to these changes. However, we must note that the observed sample is dominated by the BHBH binaries. For most models more than 90% of observed systems are double black hole binaries (Bulik & Belczyński 2003).

These results can be used as a guideline for choosing the initial conditions in numerical simulations of mergers of compact object binaries. Additionally the results of this work can be used in preparing data analysis software using templates for detection of gravitational waves from compact object inspiral. Coalescences of BHBH binaries dominate the observed sample, and we find that the observability weighted distribution is peaked around nearly equal mass binaries. We find that in most models the flux limited sample of NSNS binaries contains a large fraction of non equal mass objects. We conclude that the initial search for gravitational waves from coalescences of compact object binaries should concentrate on BHBH coalescences with mass ratio close to unity, and the low mass ratio NSNS coalescences should be taken into account.

Finally, we note that this work only includes binaries that evolved in galaxies, and neglects all possible effects, like multiple stellar interactions that are relevant for evolution in dense stellar clusters.

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REFERENCES

Abramovici, A., Althouse, W. E., Drever, R. W. P., et al. 1992, *Science*, 256, 325

Abt, H. A. 1983, *ARAA*, 21, 343

Avila-Reese, V. 1993, *Revista Mexicana de Astronomia y Astrofisica*, 25, 79

Belczyński, K. & Bulik, T. 1999, *A&A*, 346, 91

Belczynski, K. & Bulik, T. 2002, *ApJ*, 574, L147

Belczynski, K., Bulik, T., & Kluźniak, W. Ł. 2002a, *ApJ*, 567, L63

Belczynski, K., Bulik, T., & Rudak, B. 2002b, *ApJ*, 571, 394

Belczynski, K., Kalogera, V., & Bulik, T. 2002c, *ApJ*, 572, 407

Bethe, H. A. & Brown, G. E. 1998, *ApJ*, 506, 780

Blondin, J. M. 1986, *ApJ*, 308, 755

Bonazzola, S. & Marck, J. A. 1994, *Annu. Rev. Nucl. Part. Sci.*, 45, 655

Bradaschia et al. 1990, *Nucl. Instrum and Methods*, A289, 518

Brown, G. E. 1995, *ApJ*, 440, 270

Bulik, T. & Belczyński, K. 2003, *ApJ*, 589, L37

Bulik, T., Belczyński, K., & Kalogera, V. 2003, in Proceedings of the SPIE "Gravitational-Wave Detection" (SPIE, USA), 146

Bulik, T., Belczyński, K., & Rudak, B. 2004, *A&A*, 415, 407

Bulik, T., Belczyński, K., & Zbijewski, W. 1999, *MNRAS*, 309, 629

Buras, R., Rampp, M., Janka, H.-T., & Kifonidis, K. 2003, *Physical Review Letters*, 90, 241101

Burgay, M., D'Amico, N., Possenti, A., et al. 2003, *Nature*, 426, 531

Chevalier, R. A. 1989, *ApJ*, 346, 847

—. 1993, *ApJ*, 411, L33

Cordes, J. M. & Chernoff, D. F. 1998, *ApJ*, 505, 315

Danzmann, K. & et al. 1992, in *Relativistic Gravity Research*, 184

de Donder, E. & Vanbeveren, D. 1998, *A&A*, 333, 557

Delgado, A. J. & Thomas, H.-C. 1981, *A&A*, 96, 142

Duquennoy, A. & Mayor, M. 1991, *A&A*, 248, 485

Finn, L. S. & Chernoff, D. F. 1993, *Phys. Rev. D*, 47, 2198

Flanagan, É. É. & Hughes, S. A. 1998, *Phys. Rev. D*, 57, 4535

Fryer, C., Burrows, A., & Benz, W. 1998, *ApJ*, 496, 333

Fryer, C. L., Woosley, S. E., & Hartmann, D. H. 1999, *ApJ*, 526, 152

Habets, G. M. H. J. 1987, *A&A Supp.*, 69, 183

Heggie, D. C. 1975, *MNRAS*, 173, 729

Hurley, J. R., Pols, O. R., & Tout, C. A. 2000, *MNRAS*, 315, 543

Kalogera, V., Narayan, R., Spergel, D. N., & Taylor, J. H. 2001, *ApJ*, 556, 340

Kim, C., Kalogera, V., & Lorimer, D. R. 2003, *ApJ*, 584, 985

Kroupa, P. & Weidner, C. 2003, *ApJ*, 598, 1076

Kuiper, G. P. 1935, *PASP*, 47, 15

Lipunov, V. M., Postnov, K. A., & Prokhorov, M. E. 1997a, *New Astronomy*, 2, 43

—. 1997b, *MNRAS*, 288, 245

Meurs, E. J. A. & van den Heuvel, E. P. J. 1989, *A&A*, 226, 88

Nelemans, G., Yungelson, L. R., Portegies Zwart, S. F., & Verbunt, F. 2001, *A&A*, 365, 491

Nutzman, P., Kalogera, V., Finn, L. S., Hendrickson, C., & Belczynski, K. 2004, submitted to *ApJ*, astro-ph/0402091

Paczynski, B. 1990, *ApJ*, 348, 485

Portegies Zwart, S. F. & Yungelson, L. R. 1998, *A&A*, 332, 173

Scalo, J. M. 1986, *Fundamentals of Cosmic Physics*, 11, 1

Thorsett, S. E. & Chakrabarty, D. 1999, *ApJ*, 512, 288

Tsubono, K. 1995, in *First Edoardo Amaldi Conference on Gravitational Wave Experiments*, 112

Webbink, R. F. 1984, *ApJ*, 277, 355

Woosley, S. E. 1986, in *Saas-Fee Advanced Course 16: Nucleosynthesis and Chemical Evolution*, 1

Woosley, S. E., Langer, N., & Weaver, T. A. 1995, *ApJ*, 448, 315

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